

COMPARING SINGLE-EPOCH VIRIAL BLACK HOLE MASS ESTIMATORS FOR LUMINOUS QUASARS

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ABSTRACT

Single-epoch virial black hole (BH) mass estimators utilizing broad emission lines have been routinely applied to high-redshift quasars to estimate their BH masses. Depending on the redshift, different line estimators ($H\alpha$, $H\beta$, $Mg\ II\ \lambda 2798$, $CIV\ \lambda 1549$) are often used with optical/near-infrared spectroscopy. Here we use a homogeneous sample of 60 intermediate-redshift ($z \sim 1.5 - 2.2$) SDSS quasars with optical and near-infrared spectra covering CIV through $H\alpha$ to investigate the consistency between different single-epoch virial BH mass estimators. We critically compare restframe UV line estimators ($CIV\ \lambda 1549$, $CIII]\ \lambda 1908$ and $Mg\ II\ \lambda 2798$) with optical estimators ($H\beta$ and $H\alpha$) in terms of correlations between line widths and between continuum/line luminosities, for the high-luminosity regime ($L_{5100} > 10^{45.4} \text{ erg s}^{-1}$) probed by our sample. The continuum luminosities of L_{1350} and L_{3000} , and the broad line luminosities are well correlated with L_{5100} , reflecting the homogeneity of quasar spectra in the restframe UV-optical, among which L_{1350} and the line luminosities for CIV and $CIII]$ have the largest scatter in the correlation with L_{5100} . We found that the $Mg\ II$ FWHM correlates well with the FWHMs of the Balmer lines, and that the $Mg\ II$ line estimator can be calibrated to yield consistent virial mass estimates with those based on the $H\beta/H\alpha$ estimators, thus extending earlier results on less luminous objects. The CIV FWHM is poorly correlated with the Balmer line FWHMs, and the scatter between the CIV and $H\beta$ FWHMs consists of an irreducible part (~ 0.12 dex), and a part that correlates with the blueshift of the CIV centroid relative to that of $H\beta$, similar to earlier studies comparing CIV with $Mg\ II$. The $CIII]$ FWHM is found to correlate with the CIV FWHM, and hence is also poorly correlated with the $H\beta$ FWHM. While the CIV and $CIII]$ lines can be calibrated to yield consistent virial mass estimates as $H\beta$ on average, the scatter is substantially larger than $Mg\ II$, and the usage of $CIV/CIII]$ FWHM in the mass estimators does not improve the agreement with the $H\beta$ estimator. We discuss controversial claims in the literature on the correlation between CIV and $H\beta$ FWHMs, and suggest that the reported correlation is either the result based on small samples or only valid for low-luminosity objects.

Based in part on observations obtained with the 6.5 m Magellan-Baade telescope located at Las Campanas Observatory, Chile, and with the Apache Point Observatory 3.5 m telescope, which is owned and operated by the Astrophysical Research Consortium.

Subject headings: black hole physics — galaxies: active — quasars: general

1. INTRODUCTION

Knowing the mass of active supermassive black holes (SMBHs) is of fundamental importance to understanding many physical processes associated with the black hole, as well as the assembly history of the SMBH population across cosmic time. Over the past several decades, reverberation mapping (RM, e.g., Bahcall et al. 1972; Blandford & McKee 1982; Peterson 1993) has proven to be a viable technique to measure (broad-line) AGN BH mass by providing an estimate of the broad line region (BLR) size R (e.g., Peterson et al. 2004), combined with the assumptions that the BLR dynamics is dominated by the central BH mass and that the widths of the broad emission lines V are related to the virial velocity of the BLR (e.g., Dibai 1980; Wandel et al. 1999). The unknown geometry of the BLR is absorbed in a constant virial coefficient f , which is calibrated (e.g., Onken et al. 2004; Woo et al. 2010; Graham et al. 2011) to bring the products of RV^2/G into average agreement with those predicted from the local scaling relation between BH mass and bulge velocity dispersion (the $M - \sigma$ relation).

An important result of RM studies is the discovery of a tight correlation between the BLR size and the continuum luminos-

ity of broad-line AGNs (e.g., Kaspi et al. 2000; Bentz et al. 2006), i.e., the $R - L$ relation, when plotted over a wide dynamical range in AGN luminosity. This relation has led to the development of the so-called single-epoch virial BH mass estimators (“virial BH mass estimators” for short, e.g., Vestergaard 2002; McLure & Jarvis 2002; McLure & Dunlop 2004; Greene & Ho 2005; Vestergaard & Peterson 2006), in which one measures the continuum (or line) luminosity and broad line width from single-epoch spectroscopy to derive a virial product as the BH mass estimate, with coefficients calibrated from a sample of ~ 40 local AGNs with RM masses (which are further tied to the predictions from the $M - \sigma$ relation). Various versions of single-epoch virial BH mass estimators have been developed since, based on different broad lines and advocating different recipes for measuring luminosities and line widths (e.g., McGill et al. 2008; Wang et al. 2009; Vestergaard & Osmer 2009; Rafiee & Hall 2011a; Shen et al. 2011). This empirical method, albeit rooted in the RM technique, is much less expensive than RM, and hence has been applied in numerous studies to estimate quasar/AGN BH masses, notably for large statistical samples (e.g., Woo & Urry 2002; McLure & Dunlop 2004; Kollmeier et al. 2006; Greene & Ho 2007; Vestergaard et al. 2008; Shen et al. 2008, 2011).

Despite the wide application of these virial BH mass esti-

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mators, there are many statistical and systematic uncertainties of these estimates. First and foremost, all single-epoch mass estimators are bootstrapped from a sample of only ~ 40 $z \lesssim 0.4$ RM AGNs (consisting of Seyfert 1 galaxies and several PG quasars), which is known to be unrepresentative of their high-luminosity and high-redshift counterparts (e.g., Richards et al. 2011). The statistics of RM AGNs need to be substantially improved to account for the diversity in BLR properties. Secondly, different versions of virial mass estimators have different systematics depending on the quality of the spectrum and the profile of the broad line, and there is currently no consensus as to which version is the best. Nevertheless, there are some general considerations on various estimators:

- *Which line to use:* The commonly utilized pairs of line and luminosity in the restframe UV and optical are: $H\alpha$ with $L_{H\alpha}$ or L_{5100} , $H\beta$ with L_{5100} , $Mg\ II$ with L_{3000} , and CIV with L_{1350} (or L_{1450}). Since the Balmer lines $H\alpha$ and $H\beta$ are the most studied lines in reverberation mapping and the $R-L$ relation was originally measured for the Balmer line BLR radius and L_{5100} (e.g., Kaspi et al. 2000; Bentz et al. 2006), it is reasonable to argue that the virial mass estimator based on the Balmer lines is the most reliable one. The width of the broad $H\alpha$ is well correlated with that of the broad $H\beta$ and therefore it provides a good substitution in the absence of $H\beta$ (e.g., Greene & Ho 2005).

The $Mg\ II$ line has not been studied much in RM (cf., Woo 2008), and only in very few cases has a time-lag of $Mg\ II$ been measured with RM (e.g., Reichert et al. 1994; Metzroth et al. 2006); but the width of $Mg\ II$ is shown to correlate with that of $H\beta$ in single-epoch spectra (e.g., Salviander et al. 2007; Shen et al. 2008; McGill et al. 2008; Wang et al. 2009), suggesting that $Mg\ II$ may be used as a substitution for $H\beta$ in estimating virial BH masses.

The CIV line is known to vary and time-lags have been measured for CIV in several objects (e.g., Peterson et al. 2004; Kaspi et al. 2007), although the sample is too small to derive a reliable $R-L$ relation for CIV . However, the high-ionization CIV line differs from low-ionization lines such as $Mg\ II$ and the Balmer lines in many ways (for a review, see Sulentic et al. 2000), most notably it shows a prominent blueshift with respect to the low-ionization lines (e.g., Gaskell 1982). In addition, the CIV line is generally more asymmetric than $Mg\ II$ and the Balmer lines, and the width of CIV is poorly correlated with those of $Mg\ II$ and $H\beta$ (e.g., Baskin & Laor 2005; Netzer et al. 2007; Shen et al. 2008). The different properties of CIV suggest that CIV is probably more affected by a non-virial component such as arising from a radiatively-driven disk wind (e.g., Murray et al. 1995; Proga et al. 2000), and would therefore be a biased virial mass estimator (e.g., Baskin & Laor 2005; Sulentic et al. 2007; Netzer et al. 2007; Shen et al. 2008; Marziani & Sulentic 2011, and references therein). However, since both the Balmer lines and $Mg\ II$ move out of the optical bandpass at $z \gtrsim 2$, it would be useful to improve the CIV estimator in order to measure BH masses at high redshift without the need for near-IR spectroscopy.

Shen et al. (2008) used a large sample of ~ 5000 SDSS quasars to show that the difference between the CIV and $Mg\ II$ virial masses is correlated with the CIV - $Mg\ II$ blueshift. On the other hand, Assef et al. (2011) used a sample of ~ 10 quasars with optical spectra covering CIV and near-IR spectra covering $H\beta/H\alpha$ to show that the difference between CIV and Balmer line virial masses is largely driven by their rest-

frame UV-to-optical continuum luminosity ratio L_{1350}/L_{5100} , suggesting that much of the dispersion in their virial mass difference is caused by the poor correlation between L_{5100} and L_{1350} rather than between their line widths. A larger sample is needed to test this result.

- *Line dispersion vs FWHM:* The two common choices of line width are FWHM, and the second moment of the line (line dispersion, σ_{line}). Both FWHM and σ_{line} have advantages and disadvantages. FWHM is easier to measure, less susceptible to noise in the wings and line blending than σ_{line} , but is more sensitive to the treatment of the narrow line removal. Arguably σ_{line} is a better surrogate for the virial velocity (e.g., Collin et al. 2006), although the evidence is not very strong. Since currently all the RM BH masses are computed using $\sigma_{\text{line,rms}}$ measured from the rms spectra (Peterson et al. 2004), ideally one would like to use σ_{line} , albeit not measured from rms spectra, in single-epoch virial mass estimators. In practice, however, σ_{line} measured from single-epoch spectra depends on the quality of the spectra, line profile, and specific treatment of deblending, and could differ significantly from one observation/analysis to another (e.g., Denney et al. 2009; Fine et al. 2010; Rafiee & Hall 2011b; Assef et al. 2011). Therefore in terms of readiness and repeatability, σ_{line} is less favorable than FWHM in single-epoch virial mass estimators. For these reasons, we will not utilize line dispersion in the current study.

In this paper we investigate the reliability of the UV virial mass estimators (in particular CIV) compared with the $H\beta$ (or $H\alpha$) estimator with a carefully selected sample of quasars with good CIV to $Mg\ II$ coverage in optical SDSS spectra and our own near-IR spectra covering $H\beta$ and $H\alpha$. Our sample probes the high-luminosity regime ($L_{5100} > 10^{45.4} \text{ erg s}^{-1}$, or $L_{\text{bol}} \gtrsim 2.5 \times 10^{46} \text{ erg s}^{-1}$) of quasars, and thus such a study will provide confidence on estimating virial BH masses for the most luminous quasars (such as $z \gtrsim 6$ quasars). Our sample is substantially larger than earlier samples in similar studies, which enables us to draw more statistically significant conclusions. We are interested in examining empirical correlations between line widths and continuum luminosities of two different lines, and any dependence of their virial mass difference on specific quasar properties. We describe our sample and follow-up near-IR observations in §2. The procedure of measuring spectral properties is detailed in §3 and the results are presented in §4. We discuss the results in §5 and conclude in §6. Throughout this paper we adopt a flat Λ CDM cosmology with $\Omega_{\Lambda} = 0.7$, $\Omega_0 = 0.3$ and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

2. DATA

2.1. Sample Selection

We select our targets from the SDSS DR7 quasar catalog (e.g., Schneider et al. 2010; Shen et al. 2011) for follow-up near-IR spectroscopy with the following two criteria:

- redshift between 1.5 and 2.2 and avoiding redshift ranges where the $H\beta$ and $H\alpha$ lines fall in the telluric absorption bands in the near-infrared;
- with good ($S/N > 10$) SDSS spectra covering CIV through $Mg\ II$, and no broad absorption features or unusual continuum shapes.

These criteria by design selects luminous quasars (bolometric luminosity $L_{\text{bol}} > \text{a few} \times 10^{46} \text{ erg s}^{-1}$) as our targets, but the resulting sample still covers a range of spectral diversities such as the line width and velocity shift of each

broad lines. In addition, host contamination is generally negligible for these objects, which greatly simplifies our model fits and interpretations. Our targets have a similar color excess $\Delta(g-i)$ distribution (where $\Delta(g-i)$ is the deviation of $g-i$ color from the median $g-i$ color at each redshift; see Richards et al. 2003) as the underlying SDSS quasars in this redshift range, but do not include any dust-reddened objects (defined as $\Delta(g-i) \gtrsim 0.3$). $\sim 10\%$ (5/49) of the targets are radio-loud ($R \equiv f_{\nu, 6\text{cm}}/f_{\nu, 2500\text{\AA}} > 10$, see Shen et al. 2011) based on the *Faint Images of the Radio Sky at Twenty-Centimeters* (FIRST, White et al. 1997) catalog, and the rest 11 targets are not in the FIRST footprint as of July 16, 2008.

2.2. Near-IR Spectroscopy

We observed our targets during 2009-2011 with TripleSpec (Wilson et al. 2004) on the ARC 3.5 m telescope, and with the Folded-port InfraRed Echellette (FIRE, Simcoe et al. 2010) on the 6.5 m Magellan-Baade telescope. Table 1 summarizes our sample and follow-up observations. Below we describe the observations and data reduction for TripleSpec and FIRE data, respectively.

2.2.1. ARC 3.5m/TripleSpec

TripleSpec (Wilson et al. 2004) is a near-IR spectrograph with simultaneous $0.95-2.46\mu\text{m}$ coverage. We observed our targets during 2009-2011 semesters. The total exposure time varied from object to object due to different target brightness and observing conditions, but is typically 1–1.5 hr. We used slits with widths of both $1.1''$ and $1.5''$ during the course of the observations, and the resulting spectral resolution is $R \sim 2500-3500$. The slit was positioned at the parallactic angle in the middle of the observation, and we performed standard ABBA dither patterns to aid sky subtraction. For each object we observed a nearby A0V star as flux and telluric standard immediately before or after observing the science target.

We reduced the Triplespec data using the IDL-based pipeline APOTriplespecTool, which is a modified version of the Spextool package developed by Michael Cushing (Cushing et al. 2004). The reduction procedures include non-linearity correction, flat-fielding, wavelength calibration using OH sky lines (calibrated to vacuum wavelength), sky subtraction using adjacent exposures at nodding slit positions, cosmic-ray rejection, optimal extraction of 1-D spectra (Horne 1986), combining individual exposures, merging multiple echelle orders, and heliocentric corrections. We used the A0V-star observations for relative flux calibration and telluric correction following the technique of Vacca et al. (2003) using the `xtellcor` routine contained in the Spextool package (Cushing et al. 2004). We tied the absolute flux calibration to the Two Micron All Sky Survey (2MASS) (Skrutskie et al. 2006) H -band magnitudes using synthetic magnitudes computed from our spectrum with the 2MASS relative spectral response curves in Cohen et al. (2003). This absolute flux calibration neglects the continuum variability between the 2MASS and (spectroscopic) SDSS epochs, which is typically at the level of ~ 0.1 mag for average SDSS quasars (e.g., Sesar et al. 2007; MacLeod et al. 2011). It also neglects possible line shape variability of quasars between the two epochs of SDSS and near-IR observations, but this variation is likely negligible ($\sigma_{\text{FWHM}} < 0.05$ dex) based on repeated spectroscopy of the same objects (e.g., Wilhite et al. 2007; Park et al. 2012).

Finally, for 14 targets we have a second observation on a different night. We combined these repeated observations us-

ing the inverse-variance weighted mean of the two observations.

2.2.2. Magellan/FIRE

FIRE (Simcoe et al. 2010) is a near-IR echelle spectrometer covering the full $0.8-2.5\mu\text{m}$ band. We observed 20 targets during the nights of April 25-26, 2011, and another two targets on the nights of July 12-13, 2011. We used the $0.6''$ slit width in Echelle mode, which offers a spectral resolution of $R \sim 6000$ (50 km s^{-1}). Typical total exposure times were 45 min per target but varied from object to object. We observed our targets at the parallactic angle, and for each target we observed a nearby A0V star for flux and telluric standard.

We reduced the FIRE data using the IDL-based pipeline “FIREHOSE” developed by Robert Simcoe et al.³. The reduction procedures are similar to those for the TripleSpec data with the exception of sky subtraction. Instead of subtracting adjacent nodding exposures, sky subtraction was performed using a B-spline model of the sky directly constructed for each exposure following the technique of Kelson (2003).

The FIRE spectra have a substantial spectral overlap with the SDSS spectra. Therefore we used the common part with the SDSS spectrum to normalize the FIRE spectral flux density. As for our TripleSpec data, we neglect variations in line shape between the SDSS and FIRE spectroscopic epochs.

3. SPECTRAL MEASUREMENTS

To derive line width and continuum luminosities used in single-epoch virial mass estimators, we perform spectral fits to the optical and near-IR spectra, as commonly adopted in the literature (e.g., Greene & Ho 2005; Salviander et al. 2007; Shen et al. 2008; Wang et al. 2009). Spectral fits with some functional form have certain advantage of being less susceptible to noise than direct spectral measurements, although sometimes there are still ambiguities in decomposing the spectrum into different components.

Here we perform least- χ^2 global fits to the combined optical and near-IR spectra for the same object. Such global fits were not possible for objects with limited wavelength coverage (e.g., Shen et al. 2011). Each combined spectrum was de-reddened for Galactic extinction using the Cardelli et al. (1989) Milky Way reddening law and $E(B-V)$ derived from the Schlegel et al. (1998) dust map. The spectrum was then shifted to restframe using the improved redshifts provided by Hewett & Wild (2010) for SDSS quasars, where the spectral fit was performed. For each object we masked out narrow absorption line features imprinted on the spectrum, which will bias the continuum and emission line fits.

3.1. Pseudo-Continuum Fit

We first fit a pseudo-continuum model to account for the power-law (PL) continuum, Fe II emission and Balmer continuum underneath the broad emission lines of interest. All components were fit simultaneously. Templates for Fe II and Fe III emission have been constructed from the spectrum of the narrow-line Seyfert 1 galaxy, I Zw 1 (e.g., Boroson & Green 1992; Vestergaard & Wilkes 2001; Tsuzuki et al. 2006). In this work we do not include additional Fe III emission in the fits as we found this component is poorly constrained (e.g., Greene et al. 2010). For the UV Fe II template, we use the Vestergaard & Wilkes (2001) template

³ http://web.mit.edu/~rsimcoe/www/FIRE/ob_data.htm

TABLE 1
SAMPLE SUMMARY

Object Name (1)	RA (J2000) (2)	DEC (J2000) (3)	Plate (4)	Fiber (5)	MJD (6)	z_{HW} (7)	i_{PSF} (8)	$J_{2\text{MASS}}$ (9)	$H_{2\text{MASS}}$ (10)	$K_{s,2\text{MASS}}$ (11)	NIR Obs. (12)	Obs. UT (13)
J0029-0956	00 29 48.04	-09 56 39.4	0653	640	52145	1.618	17.672	16.728	15.747	15.622	TSPEC	100102/101128
J0041-0947	00 41 49.64	-09 47 05.0	0655	172	52162	1.629	16.966	16.201	15.680	15.535	TSPEC	100102/101128
J0147+1332	01 47 05.42	+13 32 10.0	0429	145	51820	1.595	17.115	16.194	15.473	15.474	TSPEC	090909/091107
J0149+1501	01 49 44.43	+15 01 06.6	0429	575	51820	2.073	17.275	16.565	15.998	15.243	TSPEC	090909/101128
J0157-0048	01 57 33.87	-00 48 24.4	0403	213	51871	1.551	18.164	16.797	16.553	0.000	TSPEC	091107/101128
J0200+1223	02 00 44.50	+12 23 19.1	0427	219	51900	1.654	17.811	16.573	16.125	0.000	TSPEC	100102/101128
J0358-0540	03 58 56.73	-05 40 23.4	0464	499	51908	1.506	18.258	17.572	16.297	0.000	TSPEC	100102/101128
J0412-0612	04 12 55.16	-06 12 10.3	0465	037	51910	1.691	17.322	16.306	16.077	15.353	TSPEC	100102/101128
J0740+2814	07 40 29.82	+28 14 58.5	0888	545	52339	1.545	17.445	16.426	15.689	15.482	TSPEC	091108
J0812+0757	08 12 27.19	+07 57 32.9	2570	026	54081	1.574	17.404	16.658	15.995	16.031	TSPEC	101202
J0813+2545	08 13 31.28	+25 45 03.0	1266	219	52709	1.513	15.385	14.085	13.271	13.056	TSPEC	091108
J0813+1522	08 13 44.15	+15 22 21.5	2270	439	53714	1.545	17.541	16.472	15.805	0.000	TSPEC	101122
J0821+5712	08 21 46.22	+57 12 26.0	1872	615	53386	1.546	16.868	15.943	15.027	15.031	TSPEC	091108/100104
J0838+2611	08 38 50.15	+26 11 05.4	1930	492	53347	1.618	16.098	15.211	14.424	14.288	TSPEC	091108
J0844+2826	08 44 51.91	+28 26 07.5	1588	179	52965	1.574	18.006	17.026	16.147	15.798	TSPEC	101202
J0855+0029	08 55 43.26	+00 29 08.5	0468	111	51912	1.525	17.952	16.829	16.545	15.668	FIRE	110426
J0917+0436	09 17 54.44	+04 36 52.1	0991	284	52707	1.587	18.543	0.000	0.000	0.000	FIRE	110427
J0933+1413	09 33 18.49	+14 13 40.1	2580	347	54092	1.561	17.465	16.520	15.540	15.448	TSPEC	100126
J0941+0443	09 41 26.49	+04 43 28.7	0570	379	52266	1.567	17.824	16.954	16.084	15.706	FIRE	110427
J0949+1751	09 49 13.05	+17 51 55.9	2370	184	53764	1.675	17.143	16.137	15.626	15.332	TSPEC	100126
J1004+4231	10 04 01.27	+42 31 23.1	1217	573	52672	1.666	16.764	15.795	15.376	15.080	TSPEC	100104
J1009+0230	10 09 30.51	+02 30 52.4	0502	429	51957	1.557	18.556	17.310	0.000	0.000	FIRE	110426
J1014+5213	10 14 47.54	+52 13 20.2	0904	259	52381	1.552	17.334	16.705	15.929	15.752	TSPEC	110124
J1015+1230	10 15 04.75	+12 30 22.2	1745	148	53061	1.703	17.400	16.374	15.989	15.658	TSPEC	110124
J1046+1128	10 46 03.22	+11 28 28.1	1601	193	53115	1.607	17.784	0.000	0.000	0.000	FIRE	110426
J1049+1432	10 49 10.31	+14 32 27.1	1749	571	53357	1.540	17.740	16.813	15.590	15.458	TSPEC	100126
J1059+0909	10 59 51.05	+09 09 05.7	1220	231	52723	1.690	16.771	15.620	15.094	14.411	TSPEC	110124
J1102+3947	11 02 40.16	+39 47 30.1	1437	205	53046	1.664	17.605	16.563	16.153	15.811	TSPEC	110222
J1119+2332	11 19 49.30	+23 32 49.1	2493	077	54115	1.626	17.338	16.230	15.465	15.322	TSPEC	110124
J1125+0001	11 25 42.29	+00 01 01.3	0280	077	51612	1.692	17.305	16.503	15.573	15.137	FIRE	110427
J1138+0401	11 38 29.33	+04 01 01.0	0838	241	52378	1.567	16.887	16.064	15.169	15.426	FIRE	110426
J1140+3016	11 40 23.40	+30 16 51.5	2220	577	53795	1.599	16.680	15.827	14.903	14.989	TSPEC	100126
J1220+0004	12 20 39.45	+00 04 27.6	0288	516	52000	2.048	17.200	16.337	15.992	15.102	FIRE	110427
J1233+0313	12 33 55.21	+03 13 27.6	0520	536	52288	1.528	17.814	16.745	16.294	0.000	FIRE	110427
J1234+0521	12 34 42.16	+05 21 26.7	0846	341	52407	1.550	16.992	16.372	15.448	15.408	TSPEC	110513
J1240+4740	12 40 06.70	+47 40 03.3	1455	424	53089	1.561	17.507	16.573	15.791	0.000	TSPEC	110222
J1251+0807	12 51 40.82	+08 07 18.4	1792	427	54270	1.607	16.907	15.975	15.068	14.717	FIRE	110426
J1333+0058	13 33 21.90	+00 58 24.3	0298	455	51955	1.511	17.776	16.888	16.039	15.683	FIRE	110426
J1350+2652	13 50 23.68	+26 52 43.1	2114	105	53848	1.624	17.042	16.110	15.490	15.548	TSPEC	110222
J1354+3016	13 54 39.70	+30 16 49.2	2116	486	53854	1.553	17.680	17.137	15.574	15.844	TSPEC	110422
J1419+0606	14 19 49.39	+06 06 54.0	1826	183	53499	1.649	17.176	16.661	15.935	0.000	FIRE	110426
J1421+2241	14 21 08.71	+22 41 17.4	2786	589	54540	2.188	16.906	15.632	14.962	14.019	TSPEC	100520/110513
J1428+5925	14 28 41.97	+59 25 52.0	0789	591	52342	1.660	17.418	16.800	15.803	15.461	TSPEC	110414/110418
J1431+0535	14 31 48.09	+05 35 58.0	1828	300	53504	2.095	16.523	15.368	14.892	14.166	TSPEC	100520
J1432+0124	14 32 30.57	+01 24 35.1	0535	054	51999	1.542	17.640	16.406	15.900	15.525	FIRE	110427
J1436+6336	14 36 45.80	+63 36 37.9	2947	444	54533	2.066	16.528	15.443	15.014	14.201	TSPEC	100520/110513
J1521+4705	15 21 11.86	+47 05 39.1	1331	256	52766	1.517	17.531	16.668	15.836	0.000	TSPEC	110422
J1538+0537	15 38 59.45	+05 37 05.3	1836	377	54567	1.684	17.889	16.905	16.179	0.000	FIRE	110426
J1542+1112	15 42 12.90	+11 12 26.7	2516	165	54240	1.540	17.295	17.083	15.702	0.000	FIRE	110427
J1552+1948	15 52 40.40	+19 48 16.7	2172	390	54230	1.613	17.450	16.547	15.903	15.746	TSPEC	110414/110418
J1604-0019	16 04 56.14	-00 19 07.1	0344	155	51693	1.636	17.072	16.219	15.281	15.420	FIRE	110426
J1621+0029	16 21 03.98	+00 29 05.8	0364	353	52000	1.689	18.489	17.255	0.000	0.000	FIRE	110714
J1710+6023	17 10 30.20	+60 23 47.5	0351	004	51780	1.549	17.359	16.446	15.426	15.092	TSPEC	110414/110418
J2040-0654	20 40 09.62	-06 54 02.5	0634	088	52164	1.611	18.850	0.000	0.000	0.000	FIRE	110426
J2045-0101	20 45 36.56	-01 01 47.9	0982	278	52466	1.661	16.415	15.650	14.889	14.672	FIRE	110427
J2045-0051	20 45 38.96	-00 51 15.5	0982	277	52466	1.590	18.063	17.047	0.000	0.000	FIRE	110426
J2055+0043	20 55 54.08	+00 43 11.4	0984	326	52442	1.624	18.594	17.308	0.000	0.000	FIRE	110427
J2137+0012	21 37 48.44	+00 12 20.0	0989	585	52468	1.670	18.046	16.864	16.096	16.042	FIRE	110713
J2232+1347	22 32 46.80	+13 47 02.0	0738	520	52521	1.557	17.340	16.208	15.531	15.372	TSPEC	091107
J2258-0841	22 58 00.02	-08 41 43.7	0724	571	52254	1.496	17.459	16.893	16.288	0.000	TSPEC	091107

NOTE. — Summary of the sample of SDSS quasars for which we have conducted near-infrared spectroscopy. Columns (4)–(6): plate, fiber and MJD of the optical SDSS spectrum for each object; (7): improved quasar redshift from Hewett & Wild (2010); (8): SDSS i -band PSF magnitudes; (9)–(11): 2MASS (Vega) magnitudes; (12): instrument for the near-IR spectroscopy; (13): UT dates of the near-IR observations. Note that here the 2MASS magnitudes were taken from Schneider et al. (2010), where aperture photometry was performed upon 2MASS images to detect faint objects, hence these near infrared data go beyond the 2MASS All-Sky and “6×” point source catalogs (see Schneider et al. 2010, for details).

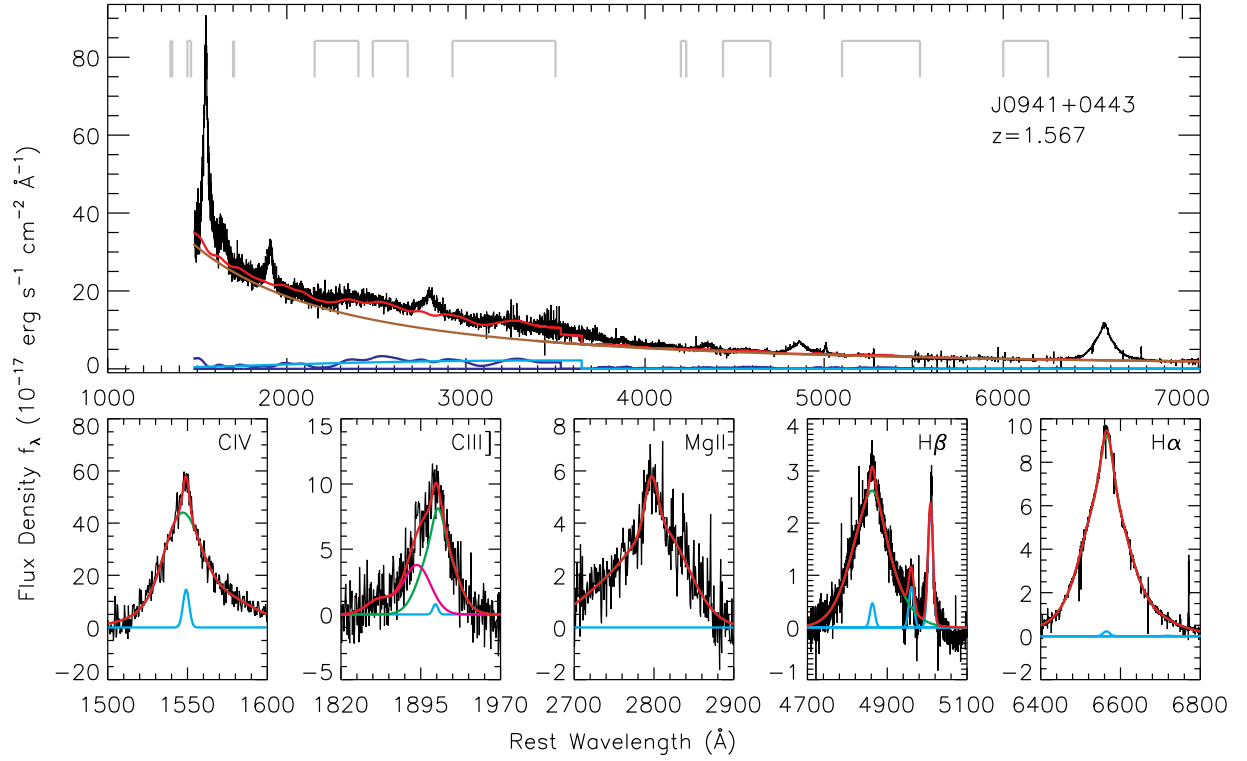


FIG. 1.— An example of our model fits to the combined optical and near-IR spectra (J0941+0443). The top panel shows the global fit of the pseudo-continuum, where the brown line is the power-law continuum, the blue line is the Fe II template fit, the cyan line is the Balmer continuum model, and the red line is the combined pseudo-continuum model to be subtracted off. The line segments near the top indicate the wavelength windows used for the pseudo-continuum fit. The bottom panels show the emission line fits to CIV through H α , where the cyan lines are the model narrow line emission, the green lines are the model broad line emission, and the red lines are the combined model line profiles. For CIII] we also show the modeled AlIII and SiIII emission in magenta.

(1000-3090 Å). Salvander et al. (2007) modified this template by extrapolating below the Mg II line, and we use their template for the 2200-3090 Å region; we augment the 3090-3500 Å region using the template derived by Tsuzuki et al. (2006). For the optical Fe II template (3686-7484 Å) we use the one provided by Boroson & Green (1992). The PL continuum model has two free parameters, the normalization and the PL slope. The UV and optical Fe II templates are fitted independently, each has three free parameters, the normalization factor, the velocity dispersion (to be convolved with the template), and the wavelength shift of the template. The Fe II templates are only used as an approximation to remove significant iron emission, and we found that they did a reasonably good job. However, we are not concerned with the properties of the iron emission in this work, and thus we do not interpret the physical meanings of the velocity dispersion and wavelength shift of the Fe II templates.

For the Balmer continuum we follow the empirical model by Grandi (1982) as composed of partially optically thick clouds with an effective temperature (e.g., Dietrich et al. 2002; Wang et al. 2009; Greene et al. 2010):

$$f_{\text{BC}}(\lambda) = AB_{\lambda}(\lambda, T_e)(1 - e^{-\tau_{\lambda}}); \quad \lambda \leq \lambda_{\text{BE}} \quad (1)$$

where T_e is the effective temperature, $\lambda_{\text{BE}} \equiv 3646 \text{ Å}$ is the Balmer edge, $\tau_{\lambda} = \tau_{\text{BE}}(\lambda/\lambda_{\text{BE}})^3$ is the optical depth with τ_{BE} the optical depth at λ_{BE} , A is the normalization factor, and $B_{\lambda}(\lambda, T_e)$ is the Planck function at temperature T_e . During the continuum fits, we have three free parameters, $1 \times 10^4 < T_e < 5 \times 10^4 \text{ K}$, $0.1 < \tau_{\text{BE}} < 2$, and $A > 0$.

Note that for limited wavelength fitting range (i.e., $\lambda < 3646 \text{ Å}$), the Balmer continuum cannot be well constrained and is degenerate with the power-law and Fe II components (e.g., Wang et al. 2009), and is generally not fitted (e.g., Shen et al. 2011). In the case of global fits, a single power-law continuum is required to simultaneously fit the region from C IV to H α , providing some additional constraints on the Balmer continuum; however even in this case, the Balmer continuum may still be poorly constrained in a few cases, which will lead to uncertainties in the power-law continuum luminosity estimates. Nevertheless, the isolation of broad emission lines is not affected much by including or excluding the Balmer continuum model.

We fit the pseudo-continuum model to a set of continuum windows free of strong emission lines (except for Fe II): 1350-1360 Å, 1445-1465 Å, 1700-1705 Å, 2155-2400 Å, 2480-2675 Å, 2925-3500 Å, 4200-4230 Å, 4435-4700 Å, 5100-5535 Å, 6000-6250 Å, 6800-7000 Å. We try fitting both with and without the Balmer continuum component and adopt the fit with the lower reduced χ^2 value; usually adding the Balmer continuum improves the global fit.

In a few cases (~ 5 objects) we found that this global pseudo-continuum model does not fit the C IV-C III] region well, which is likely caused by intrinsic reddening in these systems, ill-determined Balmer continuum strength, or mismatched iron template. For these objects we perform local ($\lambda < 2165 \text{ Å}$ with the same continuum windows defined above) continuum fits around the C IV and C III] regions without the Balmer continuum and Fe II emission (i.e., only with the PL component), in order to get better measurements for C IV and C III]. The PL continuum model is then used to estimate the monochromatic continuum luminosity $L_{\lambda} = \lambda f_{\lambda}$ at 5100 Å, 3000 Å and 1350 Å. Although some of our targets do not have spectral coverage of the restframe 1350 Å due

to their relatively lower redshift, the global PL component is well constrained so the extrapolation of the model continuum to 1350 Å is not a problem.

3.2. Emission Line Fits

Once we have constructed the pseudo-continuum model, we subtract it from the original spectrum, leaving the emission-line spectrum. We then fit the H α , H β , Mg II, C III], C IV broad line complexes simultaneously with mixtures of Gaussians (in logarithmic wavelength), as detailed below:

- H α : we fit the wavelength range 6400-6800 Å. We use up to 3 Gaussians for the broad H α component, 1 Gaussian for the narrow H α component, 2 Gaussians for the [N II] $\lambda 6548$ and [N II] $\lambda 6584$ narrow lines, and 2 Gaussians for the [S II] $\lambda 6717$ and [S II] $\lambda 6731$ narrow lines. Since the narrow [N II] lines are underneath the broad H α profile, we tie their flux ratio to be $f_{6584}/f_{6548} = 3$ to reduce ambiguities in decomposing the H α complex.
- H β : we fit the wavelength range 4700-5100 Å. We use up to 3 Gaussians for the broad H β component and 1 Gaussian for the narrow H β component. We use 2 Gaussians for the [O III] $\lambda 4959$ and [O III] $\lambda 5007$ narrow lines. Given the quality of the near-IR spectra, we decided to only fit single Gaussians to the [O III] $\lambda \lambda 4959, 5007$ lines, and we tie the flux ratio of the [O III] doublet to be $f_{5007}/f_{4959} = 3$.
- Mg II: we fit the wavelength range 2700-2900 Å. We use up to 3 Gaussians for the broad Mg II component and 1 Gaussian for the narrow Mg II component. We do not try to fit the Mg II lines as a doublet, as the line splitting is small enough not to affect the broad line width measurements, and the spectral quality is usually inadequate for fitting a doublet to the narrow Mg II emission.
- C III]: we fit the wavelength range 1820-1970 Å. We use up to 2 Gaussians for the broad C III] component and 1 Gaussian for the narrow C III] component. We use two additional Gaussians for the Si III] $\lambda 1892$ and Al III] $\lambda 1857$ lines adjacent to C III]. To reduce ambiguities in decomposing the C III] complex, we tie the centroids of the two Gaussians for the broad C III] i.e., the broad C III] profile is forced to be symmetric; we also tie the velocity offsets of Si III] and Al III] to their relative laboratory velocity offset.
- C IV: we fit the wavelength range 1500-1600 Å. We use up to 3 Gaussians for the broad C IV component and 1 Gaussian for the narrow C IV component. We do not fit the 1640 Å He II feature as its contribution blueward of 1600 Å is negligible and will not bias the C IV line fit.
- During the line fitting, all narrow line components are constrained to have the same velocity offset and line width. We also impose an upper limit of 1200 km s^{-1} for the FWHM of the narrow line component⁴.

⁴ This upper limit is slightly larger than the values used in some studies (typically $\sim 750 - 1000 \text{ km s}^{-1}$). For luminous SDSS quasars, [O III] FWHM values exceeding $\sim 1000 \text{ km s}^{-1}$ are often seen (e.g., Shen et al. 2011). We hereby adopt the 1200 km s^{-1} upper limit for the narrow line width.

While the presence of narrow line components for the Balmer lines is beyond doubt, the relative contribution from narrow line components for the UV lines is less certain. For Mg II, there is clear evidence that a narrow line component is present at least in some quasars (e.g., Shen et al. 2008, 2011; Wang et al. 2009). For C IV, the Vestergaard & Peterson (2006) virial mass calibration uses the FWHM from the whole line profile, while some argue that a narrow line component should be subtracted for C IV as well (e.g., Baskin & Laor 2005). The presence of narrow emission lines in the restframe optical spectra is essential to provide constraints on the narrow line contribution for the UV lines. We will measure the C IV line width both with and without narrow line subtraction and test if it is necessary to remove narrow line emission for C IV.

3.3. Measurement Uncertainties

It is important to quantify the uncertainties in our spectral measurements. The nature of the non-linear model and multi-component fits introduces ambiguities in decomposition, and the resulting uncertainties are usually larger than those estimated from the parameter co-variance matrix of the least- χ^2 fits. We estimate the uncertainties in measured spectral quantities using a Monte-Carlo approach as in Shen et al. (2011). For each object we generate 50 random realizations of mock spectra by adding Gaussian noise to the original spectrum at each pixel using the spectral error array. Technically speaking, these mock spectra are not the exact alternative realizations of the original spectrum since the errors were added twice, but they are only slightly degraded realizations and provide a good approximation to capture the wavelength-dependent noise properties. We fit each mock spectrum with the same fitting procedure described above and derive the distribution of each measured spectral quantity (such as FWHM, velocity offset, etc). We then take the semi-quartile of the 68% range of the distribution as the nominal uncertainty of the measured quantity. This approach takes into account the statistical uncertainties due to flux errors, and systematic uncertainties due to ambiguities in decomposing multiple components.

Fig. 1 shows an example of our global fits, and we tabulate the measured quantities in Table 2. Although with different fitting recipes, the continuum and emission line measurements are consistent with the measurements with SDSS spectra alone (Shen et al. 2011), and the largest discrepancy occurs for $L_{\text{MgII,broad}}$ and L_{3000} : $\log L_{3000}$ is systematically smaller by ~ 0.12 dex, and $\log L_{\text{MgII,broad}}$ is systematically larger by ~ 0.067 dex, than the measurements in Shen et al. (2011). This is largely caused by the additional Balmer continuum model in the spectral fits. For simplicity, from now on we will by default refer to the broad line component when we mention the FWHM or luminosity of a particular line unless stated otherwise.

4. RESULTS

We now proceed to examine correlations between continuum (line) luminosities and between line widths for different lines, as well as their virial products.

4.1. Luminosity Correlations

In Fig. 2 we compare different luminosities with the continuum luminosity at 5100 Å, and we list the slopes from the bisector linear regression fits using the BCES estimator (Akritas & Bershady 1996) in Table 3. Our objects all

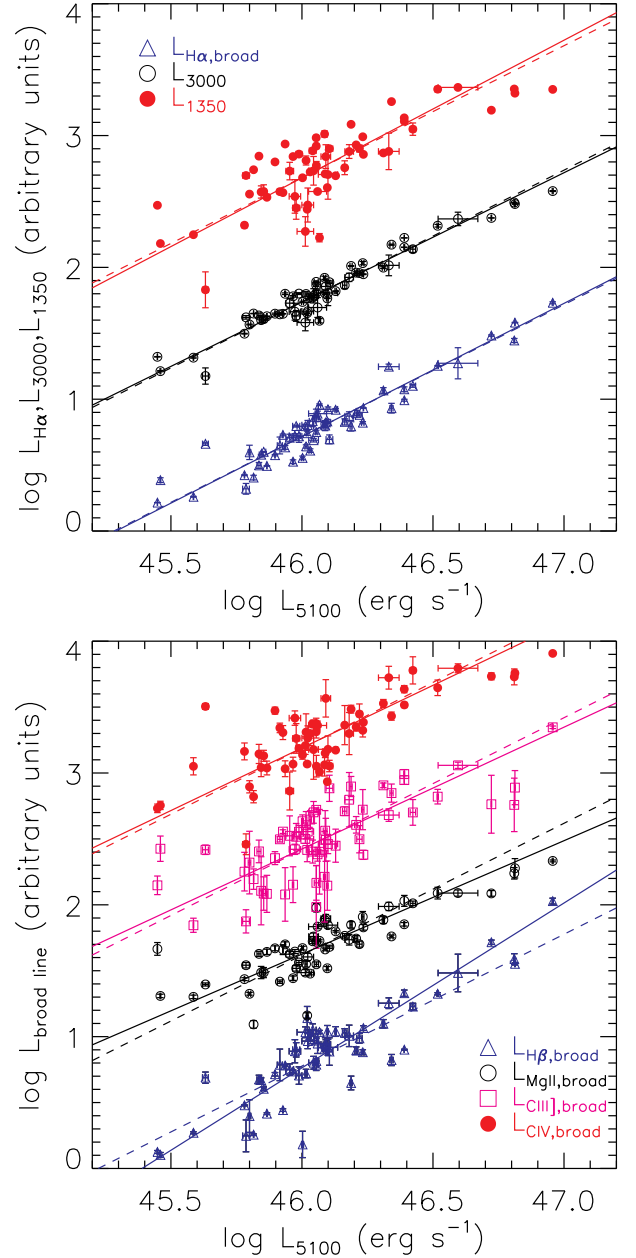


FIG. 2.— Correlations of different luminosity indicators with L_{5100} for our sample. Each correlation has been shifted vertically for clarity without changing the scatter in the correlation. *Upper*: Correlations between L_{5100} and the three most-frequently used alternative luminosity indicators for the BLR size, L_{3000} , L_{1350} and $L_{\text{H}\alpha,\text{broad}}$. *Bottom*: Correlations between L_{5100} and broad line luminosities. In both panels the solid lines are the bisector linear regression results using the BCES estimator (e.g., Akritas & Bershady 1996), and the dashed lines indicate a linear correlation of unity slope.

have luminosity $L_{5100} > 10^{45.4} \text{ erg s}^{-1}$, and therefore contamination from host starlight is mostly negligible (e.g., Shen et al. 2011). For UV estimators, L_{3000} and L_{1350} (or L_{1450}) are often used in replacement of L_{5100} . In addition, the luminosity of the broad H α line $L_{\text{H}\alpha}$ is also used in pair with H α line width (e.g., Greene & Ho 2005). Since the original $R-L$ relation is calibrated against L_{5100} , a good correlation with L_{5100} is required to produce a reasonable estimate of the BLR size with alternative luminosity indicators. As shown in the top panel of Fig. 2, all three luminosity indicators are correlated with L_{5100} , where $L_{\text{H}\alpha}$ and L_{3000} are better correlated with L_{5100}

than L_{1350} ⁵. The slopes of these correlations are very close to unity, which means the ratios of these luminosities to L_{5100} almost do not depend on luminosity over this luminosity range.

In cases where the continuum is too faint to detect or contaminated by host starlight or emission from a relativistic jet, an alternative route is to use the luminosity of the broad lines (e.g., Wu et al. 2004; Greene & Ho 2005; Shen et al. 2011). The bottom panel of Fig. 2 shows correlations of different line luminosities with L_{5100} . Again, the $H\beta$ and $Mg II$ line luminosities seem to correlate with L_{5100} with lower scatter than $C III]$ and $C IV$. Interestingly, the scatter in the $L_{C IV} - L_{5100}$ relation is comparable to that in the $L_{1350} - L_{5100}$ relation, suggesting that using the $C IV$ line luminosity will not degrade the mass estimates much than using L_{1350} . The best-fit slopes of these correlations are slightly different from unity, indicating a possible mild luminosity dependence of the ratios of these luminosities to L_{5100} over this luminosity range.

We note that these luminosity correlations are not predominantly caused by the common distance of each object. In fact, the dynamical range resulting from luminosity distances is only 0.4 dex given the limited redshift range of our objects, while the entire luminosity span is 1.5 dex. These luminosity correlations justify the usage of alternative luminosity indicators in various virial mass estimators. The scatter in the correlation between alternative luminosity indicator and L_{5100} will be one source of the scatter in virial mass estimates when compared with those based on $H\beta$ width and L_{5100} .

4.2. Line Width Correlations

While it is still debated whether or not a narrow line component for $C IV$ needs to be subtracted for high-redshift broad-line quasars (see discussions in, e.g., Bachev et al. 2004; Sulentic et al. 2007), it is clear that narrow $C IV$ emission does exist, as seen in some type 2 quasars (e.g., Stern et al. 2002). The $[O III]$ coverage in our near-IR spectra makes it possible to constrain the strength of the narrow $C IV$ emission by fixing its line width and velocity offset to those of the narrow $[O III]$ lines. We have measured the width of $C IV$ with and without the subtraction of a possible narrow line component. In Fig. 3 we compare the resulting $C IV$ FWHM with the two methods. The two objects with large error bars (J1009+0230 and J1542+1112) have associated absorption, which causes some ambiguities in decomposing the $C IV$ line in our Monte Carlo mock spectra and therefore leads to large uncertainties (see §3). We found that the narrow line contribution to $C IV$ is generally weak for objects in our sample, and only in 2 objects (J1119+2332 and J1710+6023) the narrow $C IV$ component is strong enough to make a large difference in the line width measurement. From now on we use the $C IV$ line width with narrow line subtraction. We note, however, that our quasars are luminous, and the relative strength of the narrow $C IV$ emission may be larger for lower luminosity objects (see, e.g., Bachev et al. 2004; Sulentic et al. 2007).

In the bottom panel of Fig. 3 we compare the $C IV$ FWHM and $C III]$ FWHM. The measurement errors are typically larger for $C III]$ due to the ambiguity of decomposing the $C III]$ complex, but a correlation is still seen between the FWHMs of $C IV$ and $C III]$. This is intriguing because $C III]$ does not show as large a blueshift relative to the low-ionization lines as does

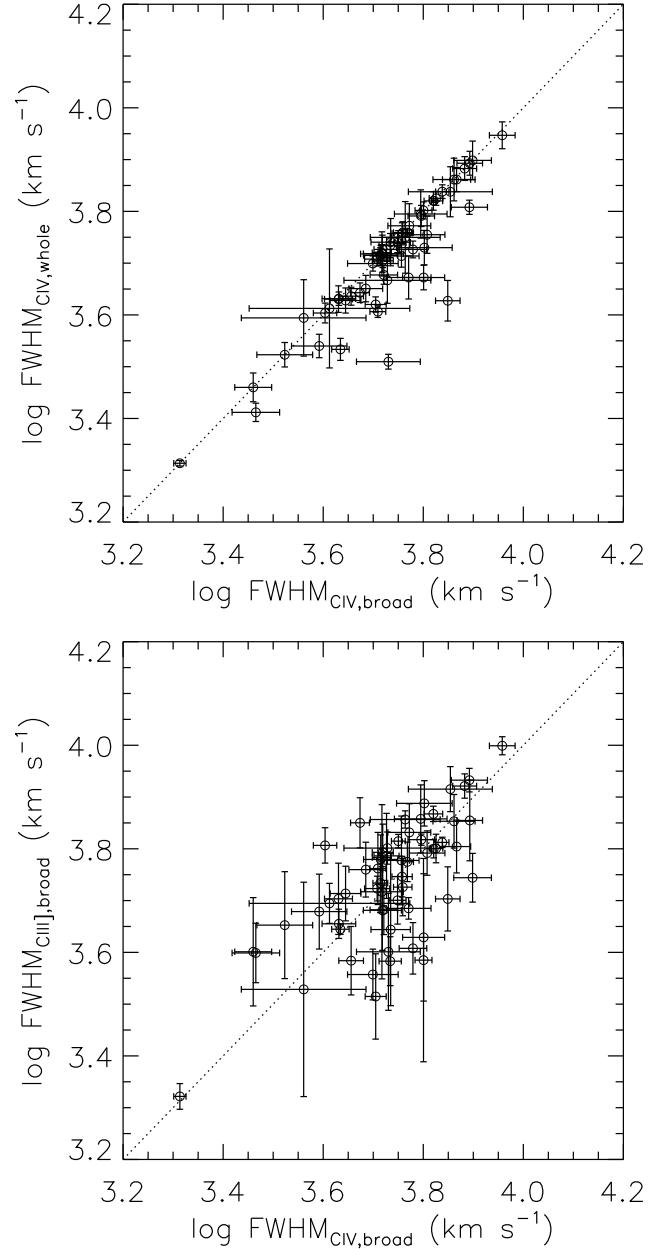


FIG. 3.— *Upper*: Comparison between the two methods of measuring the $C IV$ FWHM, i.e., with and without subtracting a narrow line component. The two methods yield similar results for the majority of our objects, indicating that the narrow line contribution is generally negligible for $C IV$ for luminous quasars with $L_{5100} > 10^{45.4} \text{ erg s}^{-1}$. *Bottom*: Comparison between the $C IV$ and $C III]$ FWHMs. The uncertainties associated with the $C III]$ FWHM measurements are typically large due to ambiguities in decomposing the $C III]$ complex, but a general correlation is seen between the two FWHMs.

$C IV$ (e.g., Richards et al. 2011, Shen et al., in preparation) when the contributions from $Si III]$ and $Al III$ are removed. In Fig. 4 we plot the FWHM against the velocity offset relative to the broad $H\beta$ line, for $C IV$ and $C III]$ respectively. In both cases a significant positive correlation is detected, i.e., FWHM increases with blueshift. Such a trend is already known when comparing $C IV$ and $Mg II$ (e.g., Shen et al. 2008, 2011), and now it is confirmed when comparing $C IV$ directly with $H\beta$. However, similar trends were not found for the FWHM of $H\alpha$, $H\beta$ or $Mg II$ against their velocity shift relative to $[O III]$. We also tested if there is any correlation between FWHM and other properties (continuum luminosity, color, line asymme-

⁵ The measured L_{3000} with the Balmer continuum component in the fit is on average smaller by ~ 0.12 dex than that without fitting the Balmer continuum. However, both measurements of L_{3000} are tightly correlated with L_{5100} .

TABLE 2
SPECTRAL MEASUREMENTS

	Format	Units	Description
objname	A10	–	Object Name
$\log L_{1350}$	F6.3	erg s^{-1}	continuum luminosity at restframe 1350 Å
Err $\log L_{1350}$	F6.3	erg s^{-1}	measurement error in $\log L_{1350}$
$\log L_{3000}$	F6.3	erg s^{-1}	continuum luminosity at restframe 3000 Å
Err $\log L_{3000}$	F6.3	erg s^{-1}	measurement error in $\log L_{3000}$
$\log L_{5100}$	F6.3	erg s^{-1}	continuum luminosity at restframe 5100 Å
Err $\log L_{5100}$	F6.3	erg s^{-1}	measurement error in $\log L_{5100}$
$\log L_{\text{CIV}}$	F6.3	erg s^{-1}	luminosity of the broad CIV line
Err $\log L_{\text{CIV}}$	F6.3	erg s^{-1}	measurement error in $\log L_{\text{CIV}}$
$\log L_{\text{CIII}}$	F6.3	erg s^{-1}	luminosity of the broad CIII] line
Err $\log L_{\text{CIII}}$	F6.3	erg s^{-1}	measurement error in $\log L_{\text{CIII}}$
$\log L_{\text{MgII}}$	F6.3	erg s^{-1}	luminosity of the broad Mg II line
Err $\log L_{\text{MgII}}$	F6.3	erg s^{-1}	measurement error in $\log L_{\text{MgII}}$
$\log L_{\text{H}\beta}$	F6.3	erg s^{-1}	luminosity of the broad H β line
Err $\log L_{\text{H}\beta}$	F6.3	erg s^{-1}	measurement error in $\log L_{\text{H}\beta}$
$\log L_{\text{H}\alpha}$	F6.3	erg s^{-1}	luminosity of the broad H α line
Err $\log L_{\text{H}\alpha}$	F6.3	erg s^{-1}	measurement error in $\log L_{\text{H}\alpha}$
FWHM_{CIV}	15	km s^{-1}	FWHM of the broad CIV line
Err FWHM_{CIV}	15	km s^{-1}	measurement error in FWHM_{CIV}
$\text{FWHM}_{\text{CIII}}$	15	km s^{-1}	FWHM of the broad CIII] line
Err $\text{FWHM}_{\text{CIII}}$	15	km s^{-1}	measurement error in $\text{FWHM}_{\text{CIII}}$
$\text{FWHM}_{\text{MgII}}$	15	km s^{-1}	FWHM of the broad Mg II line
Err $\text{FWHM}_{\text{MgII}}$	15	km s^{-1}	measurement error in $\text{FWHM}_{\text{MgII}}$
$\text{FWHM}_{\text{H}\beta}$	15	km s^{-1}	FWHM of the broad H β line
Err $\text{FWHM}_{\text{H}\beta}$	15	km s^{-1}	measurement error in $\text{FWHM}_{\text{H}\beta}$
$\text{FWHM}_{\text{H}\alpha}$	15	km s^{-1}	FWHM of the broad H α line
Err $\text{FWHM}_{\text{H}\alpha}$	15	km s^{-1}	measurement error in $\text{FWHM}_{\text{H}\alpha}$
$V_{\text{CIV-H}\beta}$	15	km s^{-1}	blueshift of the broad CIV centroid w.r.t. the broad H β centroid
Err $V_{\text{CIV-H}\beta}$	15	km s^{-1}	measurement error in $V_{\text{CIV-H}\beta}$
$V_{\text{CIV-AIII}}$	15	km s^{-1}	blueshift of the broad CIV centroid w.r.t. the broad AIII centroid
Err $V_{\text{CIV-AIII}}$	15	km s^{-1}	measurement error in $V_{\text{CIV-AIII}}$
$V_{\text{CIII-H}\beta}$	15	km s^{-1}	blueshift of the broad CIII] centroid w.r.t. the broad H β centroid
Err $V_{\text{CIII-H}\beta}$	15	km s^{-1}	measurement error in $V_{\text{CIII-H}\beta}$
$V_{\text{MgII-[OIII]}}$	15	km s^{-1}	blueshift of the broad Mg II centroid w.r.t. the narrow [O III] centroid
Err $V_{\text{MgII-[OIII]}}$	15	km s^{-1}	measurement error in $V_{\text{MgII-[OIII]}}$
$V_{\text{H}\beta-[OIII]}$	15	km s^{-1}	blueshift of the broad H β centroid w.r.t. the narrow [O III] centroid
Err $V_{\text{H}\beta-[OIII]}$	15	km s^{-1}	measurement error in $V_{\text{H}\beta-[OIII]}$
$V_{\text{H}\alpha-[OIII]}$	15	km s^{-1}	blueshift of the broad H α centroid w.r.t. the narrow [O III] centroid
Err $V_{\text{H}\alpha-[OIII]}$	15	km s^{-1}	measurement error in $V_{\text{H}\alpha-[OIII]}$
CIV AS	F4.2	–	Asymmetry parameter of the broad CIV line

NOTE. — Format of the tabulated spectral measurements. The full table is available in the online version.

TABLE 3
BCES BISECTOR SLOPES FOR
LUMINOSITY CORRELATIONS

vs L_{5100}	α	σ_α	Scatter
L_{1350}	1.044	0.099	0.13
L_{3000}	0.979	0.036	0.05
$L_{\text{H}\alpha, \text{broad}}$	1.010	0.042	0.07
$L_{\text{H}\beta, \text{broad}}$	1.251	0.067	0.11
$L_{\text{MgII}, \text{broad}}$	0.861	0.070	0.11
$L_{\text{CIII], broad}}$	0.924	0.099	0.16
$L_{\text{CIV}, \text{broad}}$	0.950	0.085	0.14

NOTE. — α and σ_α are the slope and uncertainty (1σ) in slope from the bisector linear regression fit of each luminosity against L_{5100} . The last column lists the scatter perpendicular to the best-fit linear relation (dominated by intrinsic scatter rather than measurement errors).

try) for all five lines and found none of them is significant.

In Fig. 5 we plot different line widths against the broad

H β FWHM. We have suppressed measurement errors in these plots for clarity. In addition to the traditional FWHM, we also measure the full-width-at-third-maximum (FWTM) and full-width-at-quarter-maximum (FWQM) as alternative line width indicators. Consistent with earlier studies, we see strong correlations among the widths of H α , H β and Mg II (e.g., Greene & Ho 2005; Salviander et al. 2007; Shen et al. 2008; Wang et al. 2009). On the other hand, both CIII] and CIV line widths show poor correlations with H β line width. Table 4 lists the Spearman rank-order coefficients of these correlations. Since we found using FWTM and FWQM does not improve the correlations, we will focus on FWHM from now on.

These results suggest that CIV and CIII] have different kinematics from Mg II and the Balmer lines, and possibly originate from a different region than the low-ionization lines. However, since some dispersion in the CIV and CIII] FWHM is driven by the blueshift (e.g., Fig. 4), accounting for this dependence may reduce the difference in FWHM between CIV/CIII] and the low-ionization lines. To test this, we plot the difference in FWHM $\Delta_{\text{FWHM}} = \log(\text{FWHM}_{\text{H}\beta}/\text{FWHM}_{\text{CIII], CIV}})$, as a function of the blueshift

TABLE 4
SPEARMAN TEST RESULTS

	r_s	P_{ran}
with $\text{FWHM}_{\text{H}\beta}$		
FWHM_{CIV}	0.11	0.39
$\text{FWHM}_{\text{CIII}}$	0.14	0.29
$\text{FWHM}_{\text{MgII}}$	0.64	2.8×10^{-8}
$\text{FWHM}_{\text{H}\alpha}$	0.78	1.8×10^{-13}
$\text{FWHM}_{\text{CIV,corr}}$	0.49	6.7×10^{-5}
$\text{FWHM}_{\text{CIII,corr}}$	0.36	5.2×10^{-3}
with $\frac{\text{FWHM}_{\text{CIV}}}{\text{FWHM}_{\text{H}\beta}}$		
$\log L_{1350}$	0.02	0.89
$\log L_{5100}$	-0.03	0.83
$\log(L_{1350}/L_{5100})$	0.21	0.11
EW CIV	-0.07	0.59
$V_{\text{CIV-H}\beta}$	0.59	8.9×10^{-7}
$V_{\text{CIV-AlIII}}$	0.36	5.2×10^{-3}
CIV AS	-0.28	0.028

NOTE. — r_s is the Spearman rank-order coefficient and P_{ran} is the probability of being drawn from random distributions. $\text{FWHM}_{\text{CIV,corr}}$ and $\text{FWHM}_{\text{CIII,corr}}$ are the corrected FWHMs using the linear regression fits shown in Fig. 6.

with respect to $\text{H}\beta$ for CIII] and CIV, in Fig. 6 (top panels). The green lines are the linear regression results using the Bayesian method of Kelly (2007). We then use these best-fit linear relations to correct for the observed CIII] and CIV FWHMs:

$$\log \text{FWHM}_{\text{CIII]/CIV,corr}} = \log \text{FWHM}_{\text{CIII]/CIV}} + \alpha + \beta \Delta V, \quad (2)$$

where $\Delta V = (v_{\text{off,CIII]}/v_{\text{H}\beta} - v_{\text{off,H}\beta})$ is the blueshift relative to $\text{H}\beta$, and α and β are the best-fit coefficients of the linear regression shown in green lines in the top panels of Fig. 6; $[\alpha, \beta] = [0.062, -1.87 \times 10^{-4}]$ for CIII], and $[\alpha, \beta] = [0.136, -1.57 \times 10^{-4}]$ for CIV. The “corrected” CIII] and CIV FWHMs are plotted against the broad $\text{H}\beta$ FWHMs in the bottom panels of Fig. 6. This time significant correlations are detected for both CIII] and CIV, and CIV has the most significant improvement (see Table 4 for Spearman rank-order coefficients). Nevertheless, there is still substantial scatter (~ 0.15 dex for CIII] and ~ 0.12 dex for CIV) among these correlations. This “irreducible” scatter probably again reflects the different origins of CIII] and CIV from the low-ionization lines, which makes it difficult to bring their line widths into good agreement. We will return to this point in §5.1.

Although the blueshift relative to $\text{H}\beta$ seems a viable proxy to correct the CIII] and CIV FWHM to better agree with the $\text{H}\beta$ FWHM, it is of little practical value. One would like a proxy that can be determined from regions around the CIII] or CIV line alone. We have tried to correlate Δ_{FWHM} with $\log L_{1350}$, EW CIV, asymmetry parameters of CIV (defined as $\text{AS} \equiv \ln \left(\frac{\lambda_{\text{red}}}{\lambda_0} \right) / \ln \left(\frac{\lambda_0}{\lambda_{\text{blue}}} \right)$, where λ_0 is the peak flux wavelength, and λ_{red} and λ_{blue} are the wavelengths at half peak flux from the model fits), and blueshift relative to AlIII. We found that Δ_{FWHM} is best correlated with asymmetry parameters of CIV and the blueshift relative to AlIII at the $P_{\text{ran}} < 10^{-2}$ level, although still worse than the ones against the blueshift relative to $\text{H}\beta$ (where $P_{\text{ran}} < 10^{-6}$). Using these weak correlations to correct for CIV FWHM does not seem to reduce the scatter between the CIV and $\text{H}\beta$ FWHM much.

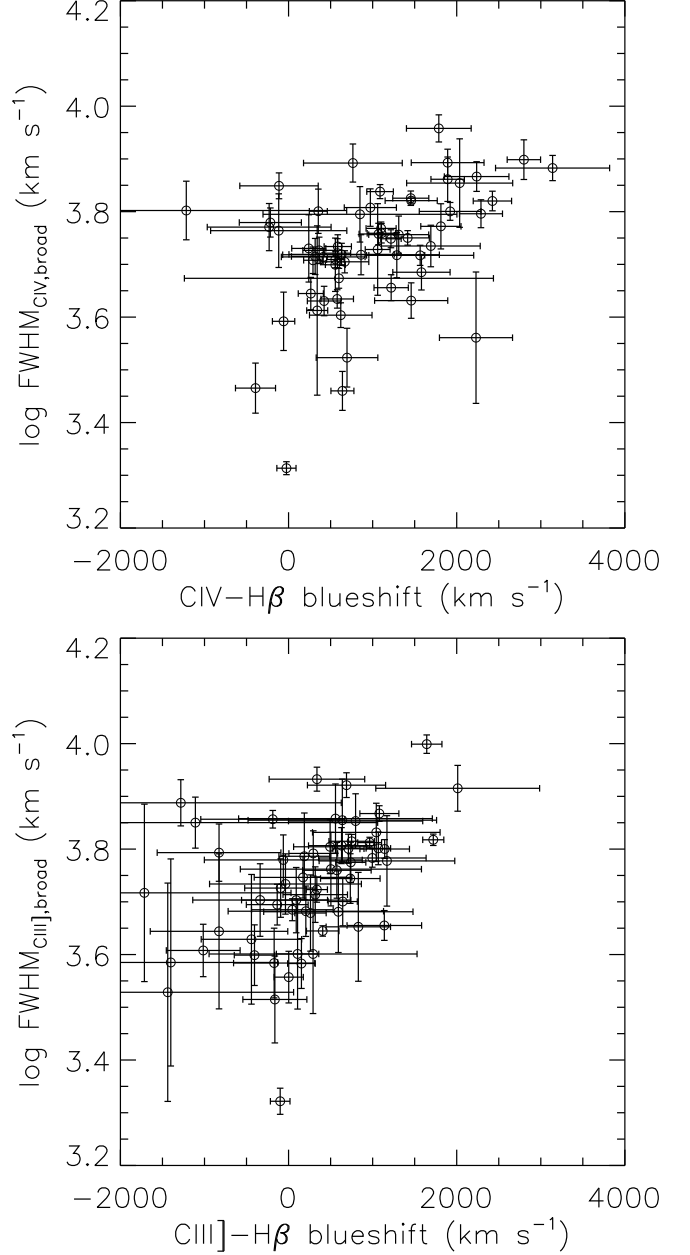


FIG. 4.— FWHM as a function of the blueshift relative to the broad $\text{H}\beta$ centroid for CIV (upper) and CIII] (bottom). Both FWHMs seem to increase with the blueshift relative to $\text{H}\beta$. Spearman correlation test results are $r_s = 0.44$ ($P_{\text{ran}} = 5.1 \times 10^{-4}$) for CIV and $r_s = 0.50$ ($P_{\text{ran}} = 4.7 \times 10^{-5}$) for CIII].

4.3. Comparing Single-Epoch Virial Mass Estimators

The investigations so far in the previous two sections treated luminosity and line width independently. In principle, if there is covariance between line width and luminosity when comparing the virial products based on different lines, the resulting scatter in the residual virial products may be increased or reduced. Since we did not observe any strong dependence of line width on luminosity for any particular line, we expect such effects to be modest at most.

As reasoned in the introduction, we adopt the $\text{H}\beta + L_{5100}$ virial masses as the “truth” values, and minimize the differences using alternative line estimator with respect to $\text{H}\beta$. There is more than one calibration based on L_{5100} and $\text{FWHM}_{\text{H}\beta}$ (e.g., McLure & Dunlop 2004; Vestergaard & Peterson 2006; Assef et al. 2011), and we use

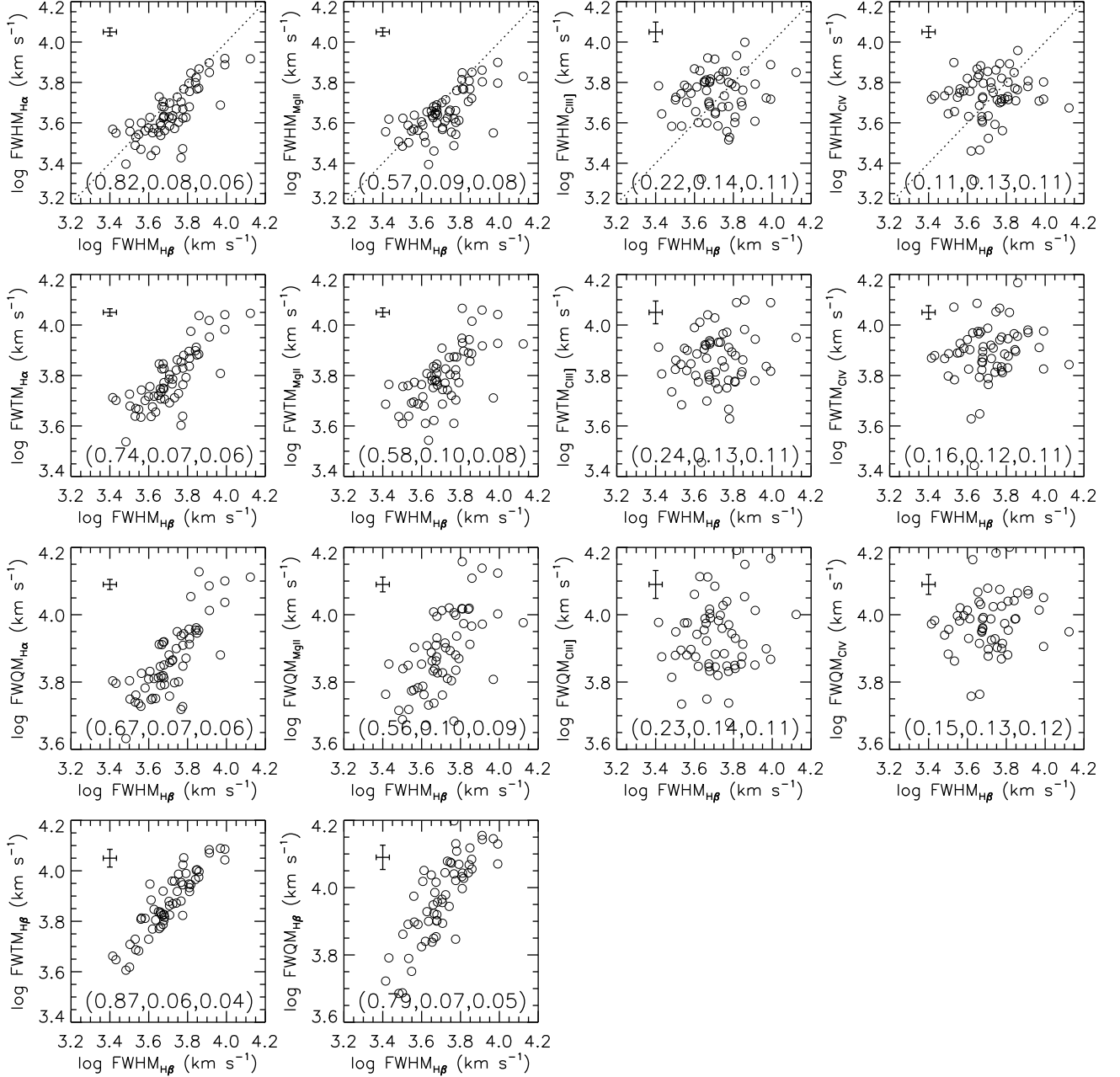


FIG. 5.— Comparisons between different line widths and $H\beta$ FWHM. Typical measurement uncertainties are indicated in each panel. The first row compares the FWHMs of $H\alpha$, $Mg\ II$, $CIII]$ and CIV with $H\beta$ FWHM, where the dotted lines show the unity relation. Only the FWHMs of $H\alpha$ and $Mg\ II$ show significant correlation with the FWHM of $H\beta$. The second and third rows show similar comparisons, but with FWTMs and FWQMs for $H\alpha$, $Mg\ II$, $CIII]$ and CIV . The conclusion is the same. The bottom two panels show the correlations between FWTM/FWQM and FWHM for $H\beta$. In each panel we show the best-fit linear regression results using the Bayesian method in Kelly (2007, predicting Y at X): the best-fit slope, the uncertainty of the slope and the intrinsic random scatter about the regression.

the Vestergaard & Peterson (2006, VP06) calibration as our standard, which is compatible with our measurements of the $H\beta$ FWHM, and provides similar estimates to those using the calibration in Assef et al. (2011, see below). The virial mass estimator based on a particular pair of line width and luminosity is:

$$\log\left(\frac{M_{BH,vir}}{M_{\odot}}\right) = a + b \log\left(\frac{L}{10^{44} \text{ erg s}^{-1}}\right) + c \log\left(\frac{FWHM}{\text{km s}^{-1}}\right), \quad (3)$$

where L and $FWHM$ are the continuum (or line) luminosity and width for the specific line, and coefficients a , b and c are

to be determined by linear regression analysis.

In order to minimize the difference in virial masses compared to our fiducial masses, we allow the slopes on both luminosity and FWHM to vary. We use the multi-dimensional Bayesian linear regression method in Kelly (2007) to perform regression, treating our standard masses as the dependent variable Y , and $(\log L, \log FWHM)$ as the 2-dimensional independent variable X . This approach takes into account measurement errors, and possible covariance between luminosity and FWHM, thus is arguably better than regressions on L versus L_{5100} and FWHM versus $FWHM_{H\beta}$ separately. The regres-

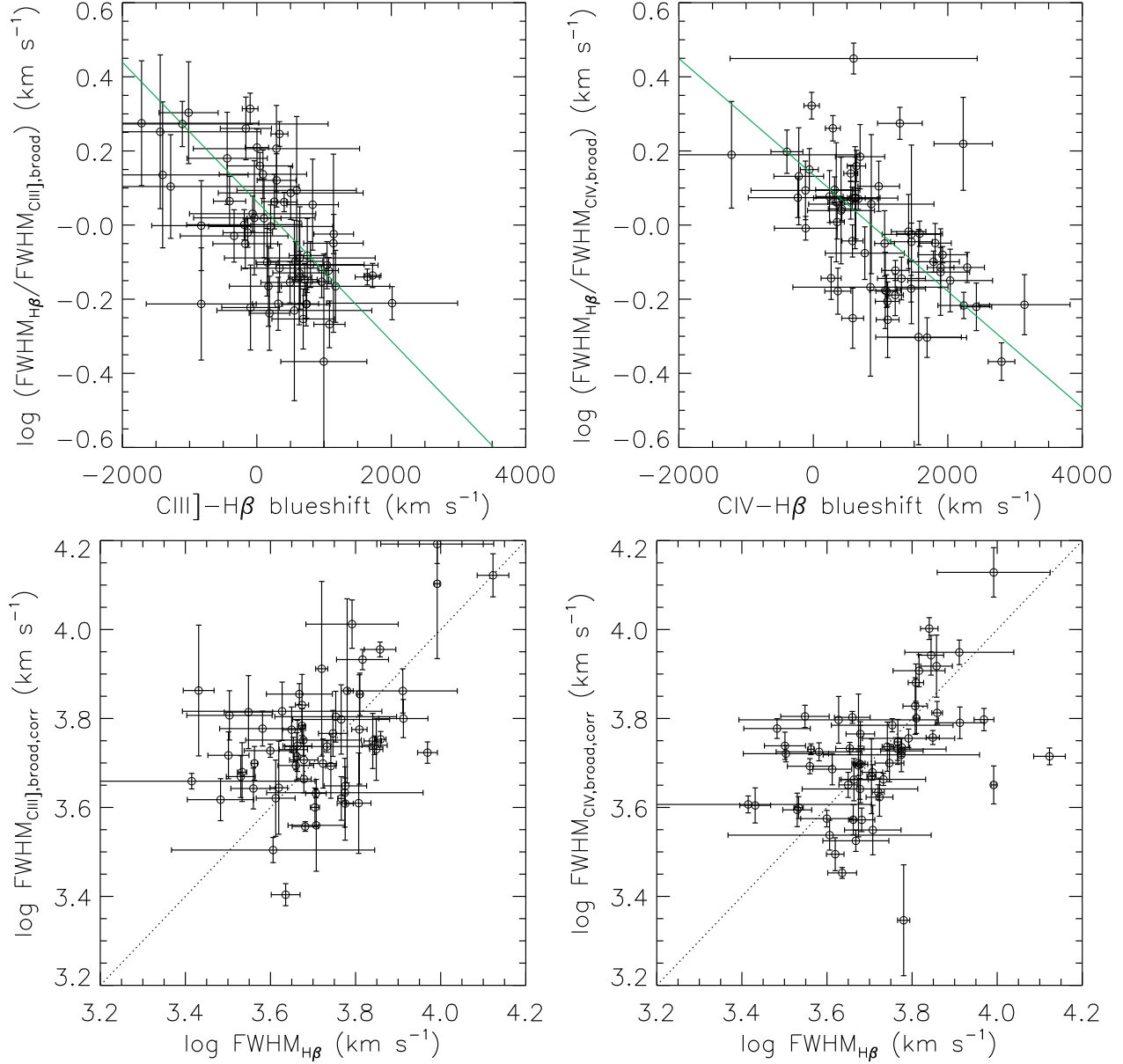


FIG. 6.— *Top*: FWHM ratio as a function of blueshift, for CIII] (left) and CIV (right), compared with H β . A significant correlation is detected for both CIII] and CIV. The green lines are the best linear regression fits using the Bayesian method in Kelly (2007). *Bottom*: “Corrected” FWHMs for CIII] (left) and CIV (right) using the best fits shown in the upper panels, compared with H β FWHM. A better one-to-one correlation is now seen between the CIII]/CIV FWHM and H β FWHM, although significant scatter still remains.

sion results are listed in Table 5.

In Fig. 7 we show the comparisons between different virial mass estimators and the VP06 H β estimator. We show in the first two columns the comparisons for several calibrations in earlier work, and in the last two columns the comparisons for our new calibrations as summarized in Table 5.

These earlier calibrations were calibrated using fainter samples than probed here, but in general they provide mass estimates that agree with the fiducial H β -based virial masses (mean offset $\lesssim 0.1$ dex). The exceptions are the Mg II-based calibrations in Shen et al. (2011), where the steeper slope of the $R-L$ relation for Mg II than for H β has led to increasingly larger discrepancies towards high luminosities⁶. For our new

calibrations, the slope on FWHM is close to (albeit slightly smaller than) 2 for H α , H β and Mg II, indicating that using FWHM in these calibrations improves the agreement with our standard mass estimator (VP06-H β). However, for CIV, our linear regression result has a slope on FWHM that is much shallower. This is because CIV FWHM is poorly correlated with H β FWHM for our sample, and the scatter between the two FWHMs rather than between the two continuum luminosities is the dominant source of the difference in their virial masses (in contrast to Assef et al. 2011, see discussions in §5.1); therefore the regression prefers a smaller dependence on CIV FWHM to minimize the difference in the two mass estimates. In other words, the individual CIV FWHM adds little to improve the agreement with our standard mass esti-

⁶ We note that the Mg II calibrations in Shen et al. (2011) were not based on linear regression fits against H β masses, and had a slope in the $R-L$ relation fixed to be the one in McLure & Dunlop (2004). Using a steeper slope 0.62 in $R-L_{3000}$ relation, the Mg II virial masses in Shen et al. (2011)

have negligible systematic offset relative to both H β -based masses at $z < 0.9$ and CIV-based masses at $z > 1.5$.

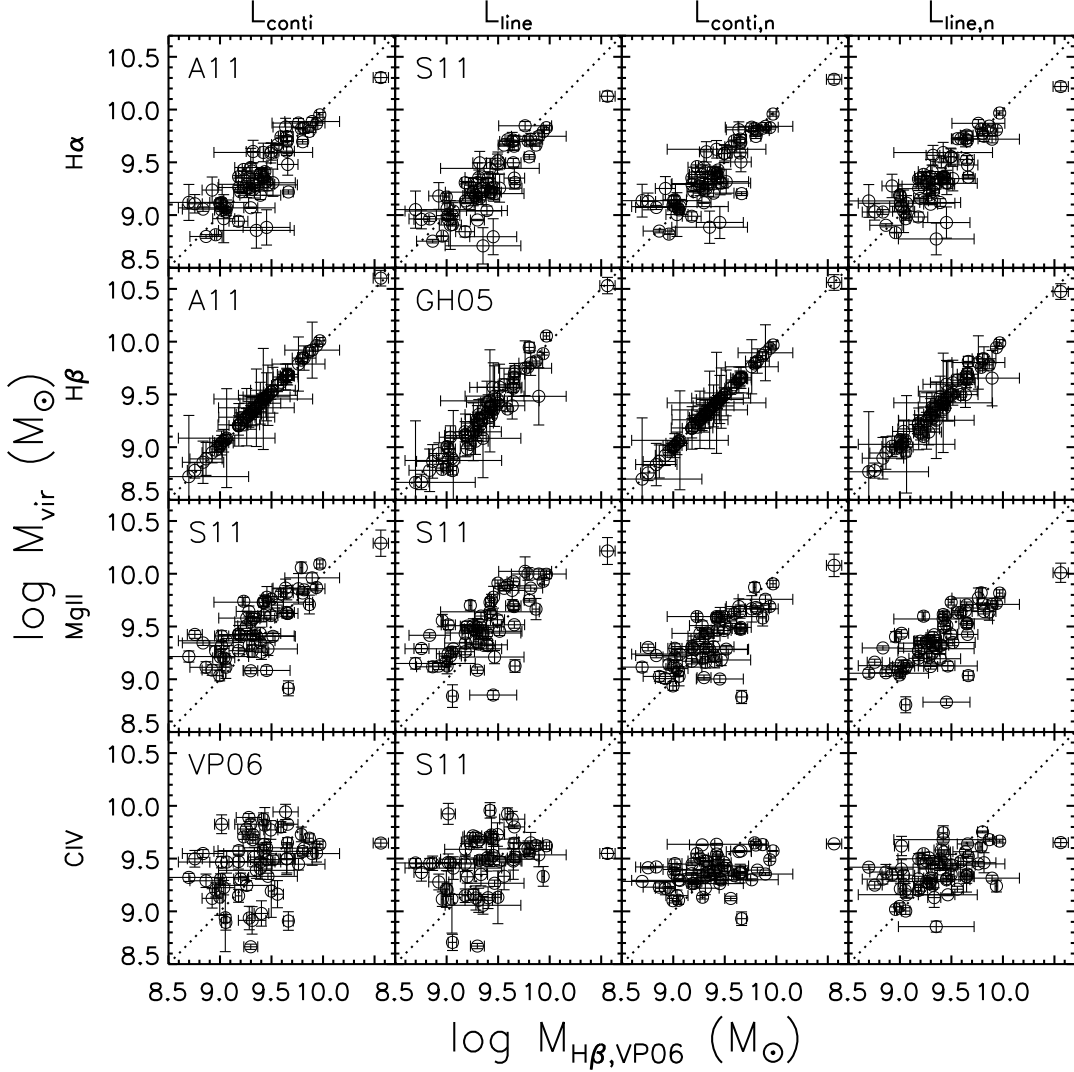


FIG. 7.— Comparisons between different virial mass estimators with the $H\beta$ estimator in Vestergaard & Peterson (2006). The left two columns show the comparisons for existing calibrations based on either continuum luminosity or line luminosity, from Assef et al. (2011, A11), Shen et al. (2011, S11), Greene & Ho (2005, GH05), and Vestergaard & Peterson (2006, VP06). The right two columns show the results for our new calibrations based on our sample, where we allow the slope on FWHM to vary for alternative lines to minimize the residual in virial mass estimates when compared to $H\beta$. The linear regression results are listed in Table 5.

mates, but instead degrades the agreement. We found similar trends for $CIII\lambda$ (not shown) as for CIV .

We test the dependence of the CIV virial mass residual on continuum luminosity, color and line shifts. The mass residual is correlated with $\log(L_{1350}/L_{5000})$, but this is expected since both masses involve luminosity. Correcting for this color dependence only marginally improves the agreement between the two masses. On the other hand, the dependence on CIV - $H\beta$ blueshift is strong enough such that incorporating this dependence can improve the agreement between CIV masses and the standard masses. These results again reflect the fact that the difference in FWHM is the dominant source in the virial mass difference. But this correction based on the CIV - $H\beta$ blueshift is of little practical use since there is no need to correct CIV -based virial masses if we have $H\beta$ coverage. Using the CIV - $AlIII$ blueshift or CIV asymmetry as a surrogate for the CIV - $H\beta$ blueshift only leads to marginal improvement of the CIV -based masses, and thus is not of much practical value either.

Our new calibrations for $MgII$ yield consistent mass estimates as those estimated from $H\beta$ for luminous ($L_{5100} >$

$10^{45.4} \text{ erg s}^{-1}$) quasars. However, it would be useful to derive a $MgII$ calibration that is also applicable to lower luminosities. For this purpose we select ~ 900 $z < 0.89$ quasars from the compilation in Shen et al. (2011) with good $H\beta$ and $MgII$ measurements ($\Delta M_{\text{vir}} < 0.1$ dex) and $L_{5100} > 10^{45} \text{ erg s}^{-1}$ (to reduce host contamination). We combine these quasars with the 60 high-luminosity quasars in our sample and perform the two-dimensional linear regression on $(\log L_{3000}, \log FWHM_{MgII})$ against the fiducial $H\beta$ -based masses. The best-fit coefficients are listed in Table 5. This $MgII$ calibration is close to the one presented in Vestergaard & Osmer (2009), and works reasonably well for the entire luminosity regime $10^{45} < L_{5100} < 10^{47} \text{ erg s}^{-1}$.

5. DISCUSSION

5.1. Comparison with Earlier Studies

There have been many studies comparing different virial mass estimators (e.g., Vestergaard & Peterson 2006; McGill et al. 2008; Dietrich & Hamann 2004; Dietrich et al. 2009; Netzer et al. 2007; Shen et al. 2008, 2011; Wang et al. 2009). These comparison studies used samples that have dif-

TABLE 5
VIRIAL MASS CALIBRATIONS

Type	a	b	c	offset (dex)	σ (dex)	Ref
Fiducial mass						
FWHM _{Hβ} , L_{5100}	0.91	0.5	2	–	–	VP06
Previous calibrations						
FWHM _{Hα} , L_{5100}	0.774	0.520	2.06	0.01	0.14	A11
FWHM _{Hα} , $L_{H\alpha}$	1.239	0.430	2.10	–0.08	0.15	S11
FWHM _{Hβ} , L_{5100}	0.895	0.520	2.00	0.03	0.003	A11
FWHM _{Hβ} , $L_{H\beta}$	1.676	0.560	2.00	–0.04	0.07	GH05
FWHM _{MgII} , L_{3000}	0.860	0.500	2.00	0.02	0.16	VO09*
FWHM _{MgII} , L_{3000}	0.740	0.620	2.00	0.17	0.17	S11*
FWHM _{MgII} , L_{MgII}	1.933	0.630	2.00	0.22	0.16	S11*
FWHM _{CIV} , L_{1350}	0.660	0.530	2.00	0.10	0.40	VP06
FWHM _{CIV} , L_{CIV}	1.525	0.457	2.00	0.09	0.38	S11
This work						
($L_{5100} > 10^{45.4} \text{ erg s}^{-1}$)						
FWHM _{Hα} , L_{5100}	1.390	0.555	1.873	0.01	0.12	–
FWHM _{Hα} , $L_{H\alpha}$	2.216	0.564	1.821	0.008	0.15	–
FWHM _{Hβ} , $L_{H\beta}$	1.963	0.401	1.959	0.02	0.04	–
FWHM _{MgII} , L_{3000}	1.816	0.584	1.712	0.02	0.16	–
FWHM _{MgII} , L_{MgII}	3.979	0.698	1.382	0.03	0.15	–
FWHM _{CIV} , L_{1350}	7.295	0.471	0.242	0.03	0.28	–
FWHM _{CIV} , L_{CIV}	7.535	0.639	0.319	0.002	0.26	–
+ $z < 0.9$ SDSS quasars						
($L_{5100} > 10^{45} \text{ erg s}^{-1}$)						
FWHM _{MgII} , L_{3000}	0.963	0.468	2.010	0.02	0.16	–

NOTE. — Virial BH mass calibrations of Eq. (3) based on different line width and luminosity combinations for the 60 objects in our sample, calibrated against the VP06 FWHM-based $H\beta$ virial masses. References of previous calibrations: GH05 (Greene & Ho 2005), VP06 (Vestergaard & Peterson 2006), VO09 (Vestergaard & Osmer 2009), S11 (Shen et al. 2011), A11 (Assef et al. 2011). For each calibration we measure the average offset and scatter σ relative to the fiducial mass $\log M_{\text{fid}}$ by fitting a Gaussian to the mass residual $\Delta = \log M - \log M_{\text{fid}}$. The scatter in the mass residual is dominated by intrinsic scatter than by measurement errors. *In using the Mg II calibrations in Shen et al. (2011) and Vestergaard & Osmer (2009) we have added 0.125 dex to the measured $\log L_{3000}$ and subtracted 0.067 dex from the measured $\log L_{MgII}$ to account for the different fitting recipe used in this work. Note that the absolute uncertainties of these calibrations are on the level of > 0.3 dex (e.g., Peterson 2011).

ferent sizes, spectral quality, luminosity and redshift ranges, and focused on different lines. The current work is among the few studies that simultaneously investigate CIV through $H\alpha$ in the same objects, and our sample is substantially larger and more homogeneous than used in similar studies (e.g., Dietrich & Hamann 2004; Dietrich et al. 2009).

Our results agree with earlier work that the line width of Mg II is well correlated with that of $H\beta$ (e.g., Salvander et al. 2007; McGill et al. 2008; Shen et al. 2008), but now we have extended this conclusion to higher luminosity than can be probed with earlier samples at lower redshift. On the other hand, we confirm the poor correlation between CIV FWHM and $H\beta$ FWHM reported earlier (e.g., Baskin & Laor 2005; Netzer et al. 2007), which was also inferred from the comparison between CIV and Mg II using SDSS quasars (e.g., Shen et al. 2008).

Assef et al. (2011) used a sample of high-redshift quasars with optical (covering CIV) and near-IR (covering the Balmer lines) spectroscopy to show that there is a correlation between the widths of CIV and the Balmer lines. They further concluded that the correlation persists, although it becomes weaker, when other objects compiled from Vestergaard & Peterson (2006), Netzer et al. (2007) and

Dietrich et al. (2009) are included.

The Assef et al. (2011) sample is small (only 9 objects with all the measurements available for Spearman tests), and they do not probe a large dynamic range in CIV blueshift (see their fig. 13). The lack of large CIV blueshift objects in their sample probably explains why they did not detect a significant correlation between the CIV virial mass residual and the CIV blueshift.

To further investigate the disagreement on the correlation between CIV and $H\beta$ FWHMs, we collected luminosity and FWHM measurements from Assef et al. (2011, 9 objects; A11), Vestergaard & Peterson (2006, 21 objects; VP06), Netzer et al. (2007, 15 objects; N07), and Dietrich et al. (2009, 9 objects; D09). We use the Prescription A measurements of CIV FWHM in Assef et al. (2011). Fig. 8 shows their distribution in the luminosity-FWHM space. The VP06 sample probes a much lower luminosity regime than the other high-redshift samples. In the bottom panel of Fig. 8 we also show the ratio of $\text{FWHM}_{\text{CIV}}/\text{FWHM}_{H\beta}$ against the continuum luminosity ratio L_{1350}/L_{5100} (optical-UV color). The A11 sample spreads over a larger range in continuum color than the other samples, including several objects that are much redder than typical quasars (e.g., Richards et al. 2003). Only the D09 and A11 samples show a mild correlation between FWHM ratio and optical-UV color, with Spearman rank-order coefficients of 0.65 ($P_{\text{ran}} = 0.02$) and 0.67 ($P_{\text{ran}} = 0.05$), respectively. This correlation helps to reduce the virial mass differences between CIV and $H\beta$ once this color effect is taken out (Assef et al. 2011).

In Fig. 9 we show the comparison between CIV and $H\beta$ FWHMs for different samples. We run Spearman tests for the combined sample and for each subsample, and find that the correlation between the CIV and $H\beta$ FWHMs reported in Assef et al. (2011) is essentially driven by objects in the VP06 sample, which probes a much fainter luminosity than other samples as indicated in Fig. 8. None of the other samples show significant correlations between the two FWHMs, and this result does not change when we restrict ourselves to high-quality measurements. These high-redshift samples have a narrower dynamic range in line width than the VP06 sample, and thus the intrinsic scatter between the CIV and $H\beta$ FWHMs can easily wash out any weak correlation. We therefore reinforced our earlier conclusion in §4.2 that, at least for the high-luminosity objects, the CIV FWHM is poorly correlated with the $H\beta$ FWHM.

5.2. Implications for High-Redshift Quasars

The comparisons between different line estimators in previous sections suggest that in the absence of Balmer lines, the Mg II estimator can be used as a substitute, which will yield consistent virial mass estimates to those based on the Balmer lines. On the other hand, CIV and CIII] can be used, although using the individually measured CIV/CIII] FWHM does not seem to offer much advantage over simply using a constant value; of course, this conclusion is valid for the high-luminosity regime probed by this study. CIV is a complicated line, and may be more affected by a non-virial component as luminosity increases (see discussions in Richards et al. 2011). These unusual properties of CIV suggest that it is likely the least reliable virial mass estimator at high-redshift, thus optical/near-IR coverage of Mg II or the Balmer lines is desired for reliable virial mass estimates (e.g., Netzer et al. 2007; Trakhtenbrot et al. 2011; Marziani & Sulentic 2011).

One should also be aware that even for the most reliable

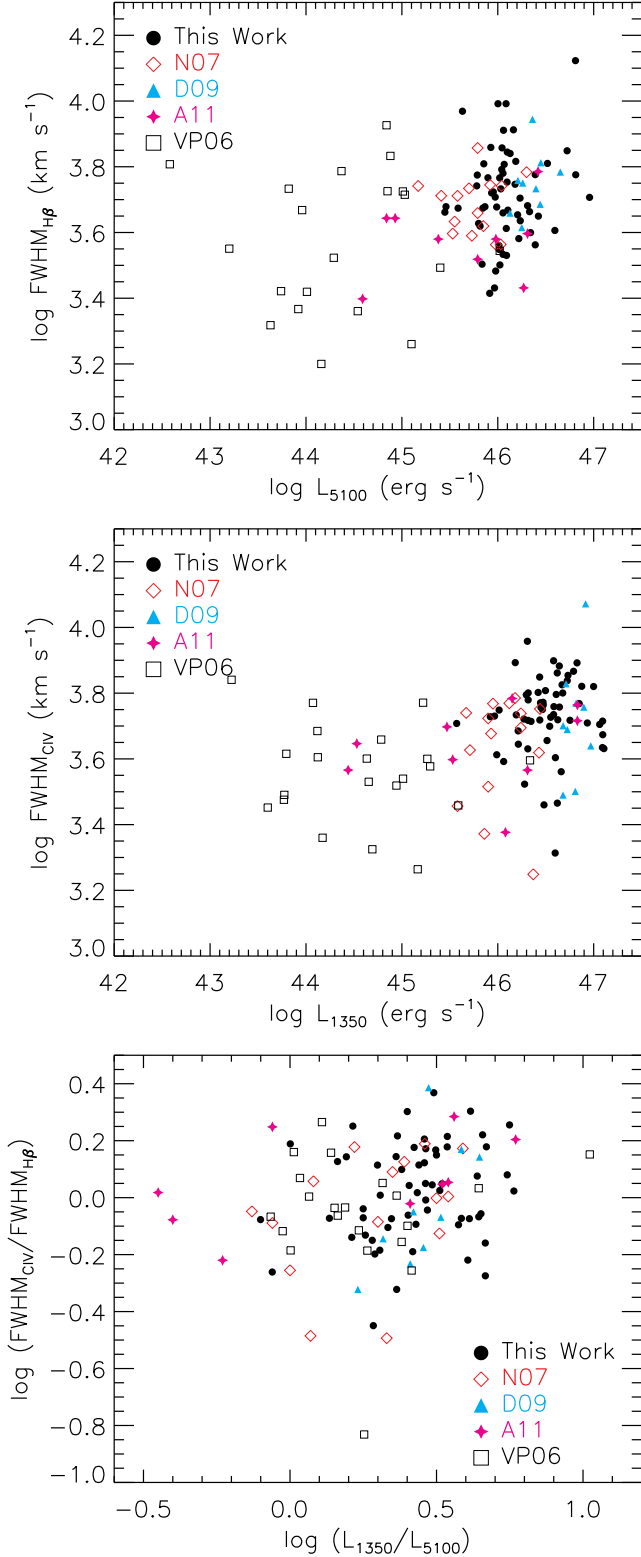


FIG. 8.— Comparisons between our sample and other samples in the literature [Assef et al. (2011, 9 objects; A11), Vestergaard & Peterson (2006, 21 objects; VP06), Netzer et al. (2007, 15 objects; N07), and Dietrich et al. (2009, 9 objects; D09)]. *Upper*: Distribution of $H\beta$ FWHM versus L_{5100} . *Middle*: Distribution of CIV FWHM versus L_{1350} . *Bottom*: Distribution of the ratio $FWHM_{CIV}/FWHM_{H\beta}$ versus continuum color L_{1350}/L_{5100} . Note that the VP06 sample has a much fainter luminosity than the other samples. This low-redshift sample also has smaller $H\beta$ and CIV FWHMs than the other high-redshift samples. Only for the samples in Dietrich et al. (2009, 9 object) and Assef et al. (2011, 9 objects) is there a significant correlation between $FWHM_{CIV}/FWHM_{H\beta}$ and L_{1350}/L_{5100} , which helps to reduce the virial mass difference between CIV and $H\beta$ once this color-dependence is taken out.

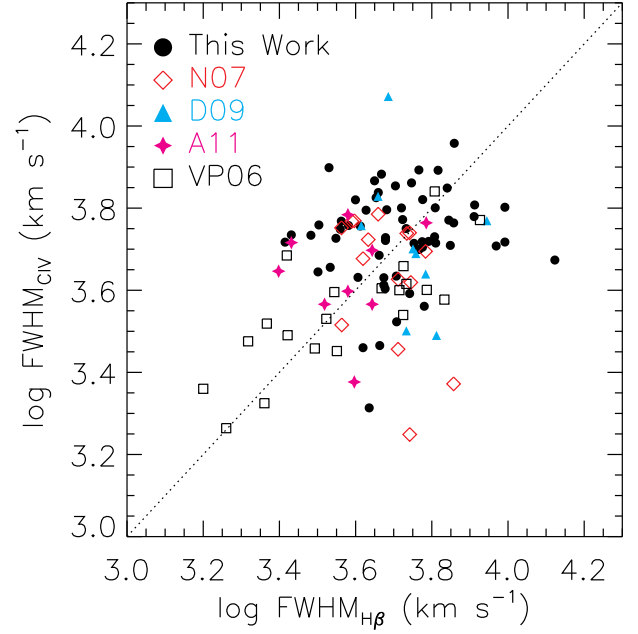


FIG. 9.— Comparison between CIV FWHM and $H\beta$ FWHM for different samples. The dotted line is the unity relation. Only for the low-redshift and low-luminosity VP06 sample is there a significant correlation between the two FWHMs.

$H\beta$ -based virial mass estimates, there is still considerable scatter between virial masses and true masses (on the level of > 0.3 dex, e.g. Peterson 2011). Rare objects with unusual continuum and/or emission line properties (for instance, dust reddened quasars, e.g., Richards et al. 2003) will lead to additional uncertainty in these virial mass estimates, along with measurement errors from poor spectral quality. Recognizing and accounting for the uncertainties in these virial mass estimates is crucial in essentially all BH mass related studies (e.g., Shen et al. 2008; Kelly et al. 2009, 2010; Shen & Kelly 2010, 2012).

6. SUMMARY

In this paper we have empirically determined the relations between single-epoch virial mass estimators based on different lines for luminous ($L_{5100} > 10^{45.4}$ erg s $^{-1}$) quasars, using a sample of 60 intermediate-redshift quasars with complete coverage from CIV through $H\alpha$ with good optical and near-IR spectroscopy. Our sample consists of typical quasars with no peculiarities in their continuum and emission line properties, has negligible contamination from host starlight, and is large enough to draw statistically significant conclusions. The main conclusions of this paper are the following:

- The Mg II FWHM is well correlated with the FWHM of the Balmer lines up to high luminosities ($L_{5100} > 10^{45.4}$ erg s $^{-1}$), which justifies the usage of Mg II (in combination with L_{3000} or L_{MgII}) to estimate virial BH masses for luminous quasars at high redshift.
- The narrow-line contribution to the CIV line is generally negligible for high-luminosity quasars; and the FWHMs of CIV and CIII] are well correlated, suggesting that both lines originate from similar regions.
- The FWHM of CIV is poorly correlated with that of the Balmer lines, suggesting different BLRs for the CIV line and for the Balmer lines. Part of the discrepancy

between the C_{IV} and H β FWHMs is correlated with the blueshift of C_{IV} relative to H β .

- Using the FWHM of C_{IV} increases the scatter between C_{IV} and H β based virial masses, which is at least true for the high luminosity regime probed in this study. C_{III]} does not seem to be a superior substitute for C_{IV}, both because of the non-correlation between C_{III]} and H β FWHMs, and because C_{III]} is blended with Si_{III]} and Al_{III}.

While in this work we focused on empirical relations, the correlations (and lack thereof) among these broad lines and continuum luminosities are ultimately determined by the accretion disk and BLR physics. In future work, we will use the same sample to investigate in more detail the emission line and continuum properties of intermediate-redshift quasars, such as the Baldwin effect (Baldwin 1977), line centroid shifts, as well as average properties and correlations therein, in order to understand the underlying physics. In the mean time, we will expand our sample to include less luminous objects and test these correlations over a larger dynamic range in quasar luminosity.

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